



Extremely Low Molecular Gas Content in a Compact, Quiescent Galaxy at $z = 1.522$

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Abstract

One of the greatest challenges to theoretical models of massive galaxy formation is the regulation of star formation at early times. The relative roles of molecular gas expulsion, depletion, and stabilization are uncertain as direct observational constraints of the gas reservoirs in quenched or quenching galaxies at high redshift are scant. We present ALMA observations of CO(2–1) in a massive ($\log M_*/M_\odot = 11.2$), recently quenched galaxy at $z = 1.522$. The optical spectrum of this object shows strong Balmer absorption lines, which implies that star formation ceased ~ 0.8 Gyr ago. We do not detect CO(2–1) line emission, placing an upper limit on the molecular H₂ gas mass of $1.1 \times 10^{10} M_\odot$. The implied gas fraction is $f_{\text{H}_2} \equiv M_{\text{H}_2}/M_* < 7\%$, $\sim 10\times$ lower than typical star-forming galaxies at similar stellar masses at this redshift, among the lowest gas fractions at this specific star formation rate at any epoch, and the most stringent constraint on the gas contents of a $z > 1$ passive galaxy to date. Our observations show that the depletion of H₂ from the interstellar medium of quenched objects can be both efficient and fairly complete, in contrast to recent claims of significant cold gas in recently quenched galaxies. We explore the variation in observed gas fractions in high- z galaxies and show that galaxies with high stellar surface density have low f_{H_2} , similar to recent correlations between specific star formation rate and stellar surface density.

Key words: galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: high-redshift – galaxies: ISM

1. Introduction

Producing realistic populations of non-star-forming or quiescent galaxies over cosmic time remains a significant challenge to current theoretical models of galaxy formation and evolution. Quenched galaxies have been identified as early as $z \sim 4$ (e.g., Straatman et al. 2014; Glazebrook et al. 2017), but their emergence peaks at $z \sim 2$, an epoch after which the majority above $\log M_*/M_\odot \gtrsim 11$ have their star formation truncated (e.g., Muzzin et al. 2013; Tomczak et al. 2014; Davidzon et al. 2017). The physical mechanisms responsible for rapidly halting the star formation in early massive galaxies, and preventing future star formation for many gigayears, remain insufficiently constrained, partly due to a poor understanding of the observable signatures of the physics affecting star formation. These processes are likely tied to either depleting, expelling, and/or heating the cold molecular gas in galaxies, which would otherwise fuel star formation (see, e.g., Man & Belli 2018, and references therein).

This has motivated investigations into the molecular gas properties of quenched galaxies, as probed by the rotational transitions of CO. The overwhelming majority of these studies have, until recently, been limited to the local universe, where observations indicate very low molecular gas fractions (0.1%–1%) and very low star formation efficiency relative to star-forming galaxies (e.g., Davis et al. 2011, 2013; Saintonge et al. 2011a, 2011b, 2012). Locally, massive quiescent galaxies have old stellar populations, indicating that their primary epoch of star formation occurred many gigayears in the past; the residual molecular gas reservoirs are likely either supplied by external

processes such as gas-rich merging (e.g., Young et al. 2014) or internally via stellar mass loss (e.g., Davis et al. 2016). Late accretion of gas will often be characterized by misaligned stellar and molecular gas kinematics, which may not be as important even as late as $z \sim 0.7$ (Hunt et al. 2018). However, recently quenched (e.g., post-starburst) galaxies, which have ceased star formation within the last $\lesssim 1$ Gyr, provide the opportunity to observe the molecular gas properties immediately following a recent star-forming episode to gain better insight into the quenching process. Observations of cold gas in local post-starburst galaxies indicate large ($\sim 10^9 M_\odot$ or $f_{\text{H}_2} \sim 0.01$ –0.3) reservoirs of cold gas despite their low star formation rates (e.g., French et al. 2015, 2018), suggesting that the recent quenching of these galaxies cannot be simply due to a lack of cold gas. These transitioning galaxies are extraordinarily rare in the local universe, $< 0.2\%$ of the overall population and rarer at high masses (e.g., French et al. 2015), but become far more common at $z > 1$ (Whitaker et al. 2012).

Exploring the molecular gas content of high-redshift transitioning galaxies is necessary to establish which physical mechanisms are responsible for building up the massive end of the red sequence. Outside of the local universe only ~ 10 $z < 1$ quiescent galaxies have published constraints on their molecular gas reservoirs based on CO lines (Suess et al. 2017; Spilker et al. 2018), indicating a large spread of gas fractions ($f_{\text{H}_2} \equiv M_{\text{H}_2}/M_*$) from upper limits of $\sim 3\%$ to measured $f_{\text{H}_2} \sim 15\%$ in galaxies below the star-forming main sequence (Spilker et al. 2018) and between 4% and 20% for massive post-starburst galaxies (Suess et al. 2017). Measurements beyond $z \sim 1$ are even more sparse, with one quiescent galaxy at $z \sim 1.4$ constrained to $f_{\text{H}_2} < 10\%$ (Sargent et al. 2015), two

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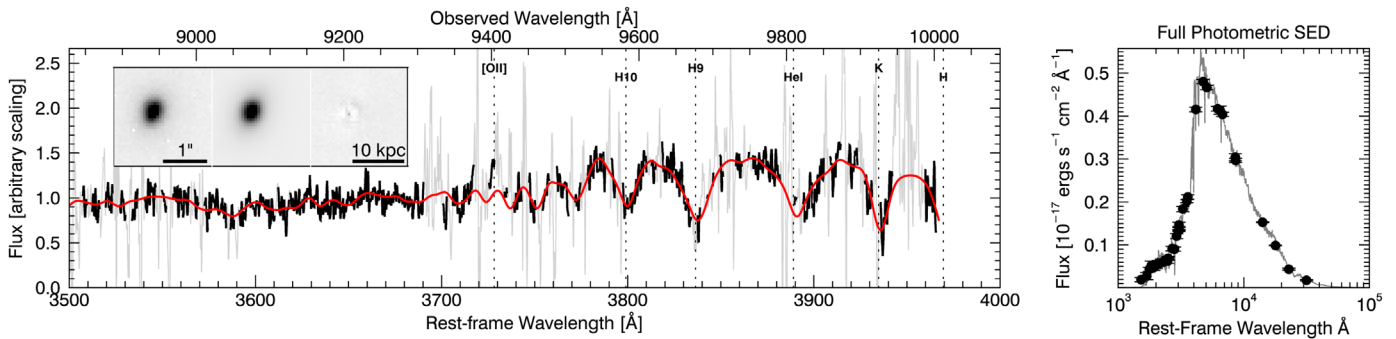


Figure 1. Left panel: Keck/LRIS optical spectrum (black) and best-fitting model (red) and inset *HST*-WFC3 F160W image, Sérsic model, and residual. Right panel: photometric spectral energy distribution (SED) for C21434 from the NMBS photometry. The strong Balmer absorption features and peaked SED reflect the young, quiescent stellar population of this galaxy.

cluster galaxies at $z = 1.46$ (Hayashi et al. 2018) and $z \sim 1.62$ (Rudnick et al. 2017) with $f_{\text{H}_2} = 35\%$ and $\sim 42\%$ respectively.⁸ In contrast with the Sargent et al. (2015) measurement, but perhaps consistent with the spread in the small number of observed individual galaxies, Gobat et al. (2018) performed a stacking analysis of dust continuum and determined that quiescent galaxies at $z \sim 1.8$ have average gas fractions $\sim 16\%$. Although these results are not in tension, there is much work to be done to define the distribution and scatter in cold gas reservoirs remaining in galaxies as they shutdown star formation at this key epoch. With this work, we quantify the molecular gas reservoir in C21434, a massive and recently quenched galaxy at $z = 1.52$, leveraging the unrivaled sensitivity of ALMA at the peak of galaxy assembly. Our ALMA observations represent the deepest constraint to date on the molecular gas content of a non-star-forming galaxy at $z > 1$. We assume a standard concordance cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$, and a Chabrier (2003) initial mass function.

2. Target Selection and Data

2.1. Multiwavelength Photometry and LRIS Spectrum

The targeted galaxy, C21434, is selected from a sample of massive galaxies ($\log M_*/M_\odot \gtrsim 11$) at $z \sim 1.5$ in the NEW-FIRM Medium Band Survey (NMBS, Whitaker et al. 2011). The sample and full analysis is described in Bezanson et al. (2013), but we briefly summarize here. The NMBS survey provides extensive multiwavelength photometry for this object from the UV to $24 \mu\text{m}$, including medium-band near-IR filters that span the Balmer/4000 Å break at the redshift of the target (shown in the right panel of Figure 1). A deep (18 hr) optical spectrum was taken in 2010 January and April using the LRIS spectrograph (main panel of Figure 1) and high-resolution rest-frame optical imaging was obtained using the F160W filter on the *HST*-WFC3 Camera (Program *HST*-GO-12167, PI: Franx). The structural parameters of C21434 are measured using *Galfit* (Peng et al. 2002) to fit a single Sérsic (1968) profile to the *HST* image. The image, best-fitting model, and residuals from the fit are included in the inset of Figure 1.

The NMBS photometry are fit using FAST (Kriek et al. 2009) with Bruzual & Charlot (2003) stellar population

synthesis models and delayed exponentially declining star formation histories, fixing to the spectroscopic redshift $z_{\text{spec}} = 1.522$. The best-fitting model finds a stellar mass of $\log M_*/M_\odot = 11.2$. The strongly peaked Balmer in the SED and strong Balmer absorption features indicate that C21434 is a recently quenched “A-type” post-starburst galaxy. The photometric fit yields a stellar age of ~ 0.8 Gyr. We also spectroscopically fit the age of C21434 using PPF (Cappellari & Emsellem 2004) to fit a linear, nonnegative sum of Vazdekis (1999) single stellar population models, again finding a stellar age of ~ 0.8 Gyr.

2.2. ALMA Observations

ALMA Band 3 observations were carried out in two consecutive observing blocks on 2016 January 18 as part of project 2015.1.00853.S (PI: Bezanson). The CO(2–1) line is redshifted to a sky frequency of 91.411 GHz. One 1.875 GHz spectral window was centered around this frequency with 7.8 MHz ($\sim 25 \text{ km s}^{-1}$) channelization, with two additional 1.875 GHz spectral windows placed for continuum observations. The total duration of the two blocks was 160 minutes, with 109 minutes on source. A total of 40 antennas were active in the array, reaching maximum baselines of 330 m, yielding a resolution of $\sim 2''$. Quasars J1058+0133 and J0949+0022 served as the bandpass and complex gain calibrators, respectively, for both observing blocks. J1058+0133 was also used to calibrate the absolute flux scale in the first observing block, while Ganymede served this purpose in the second block. We verified that the flux density of the gain calibrator was consistent between the two tracks to within 3%. The data were reduced using the standard ALMA Cycle 3 pipeline, and no significant issues with this reduction were found. The reduced data reach a continuum sensitivity of $9 \mu\text{Jy}/\text{beam}$ at 98 GHz, and a CO(2–1) line sensitivity of $105 \mu\text{Jy}/\text{beam}$ per 100 km s^{-1} channel.

2.3. Nondetection and H_2 Gas Mass Limit

To extract a spectrum of C21434, we fit a point source to the visibilities in bins of eight channels, or $\sim 200 \text{ km s}^{-1}$. The stellar effective radius, $0''.23$, is much smaller than the resolution of these observations, $\sim 2''$, and so this unresolved source is expected to be pointlike. We fix the position of the modeled point source to its position in optical/NIR imaging, leaving only the flux density at each velocity channel as a free parameter. The extracted spectrum is shown in Figure 2. We do

⁸ Using the stellar mass estimate from Skelton et al. (2014) 3D-*HST* catalogs for maximal consistency, which is ~ 0.3 dex lower than the value quoted in Rudnick et al. (2017), corresponding to $f_{\text{H}_2} = 20\%$. This discrepancy is slightly larger than the ~ 0.2 dex uncertainty expected for stellar mass estimates with a fixed IMF (Muzzin et al. 2009).

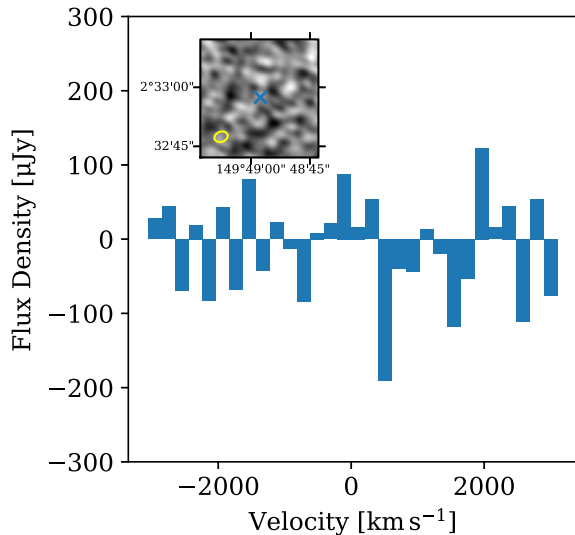


Figure 2. CO(2–1) spectrum and inset CO(2–1) map demonstrating the nondetection of C21434. The CO(2–1) map is integrated within 600 km s^{-1} of the optical redshift and the optical position of C21434 is indicated by a blue cross. We place a 3σ upper limit on the molecular gas mass of $M_{\text{H}_2} < 1.1 \times 10^{10} M_{\odot}$, corresponding to a molecular gas fraction of $< 7\%$.

not detect C21434 in either CO(2–1) emission or 3 mm continuum.

To set an upper limit on the integrated CO(2–1) flux of C21434, we assume that the gas spatial and kinematic distribution traces the stellar continuum and therefore adopt a CO line width equal to the line width determined from the stellar absorption features, which have $\text{FWHM} \sim 600 \text{ km s}^{-1}$ (Bezanson et al. 2013). This yields a 3σ upper limit to the CO(2–1) luminosity of $L'_{\text{CO}} < 2.4 \times 10^9 \text{ K km s}^{-1} \text{ pc}^2$.

In order to convert this limit to a molecular gas mass, we assume thermalized line emission, as observed in local early-type galaxies from the ATLAS^{3D} survey (Young et al. 2011). Because the CO(2–1) line is near the ground state, only a limited range of CO(2–1)/CO(1–0) excitation variations are observed in galaxies; our estimate of thermalized excitation should be accurate to $< 30\%$. We adopt a Milky Way–like CO–H₂ conversion factor, $\alpha_{\text{CO}} = 4.4 M_{\odot} / (\text{K km s}^{-1} \text{ pc}^2)$. Aside from being a conservative choice, this value is motivated by observations of local quiescent galaxies (Young et al. 2011) and theoretical models of the variations of α_{CO} with metallicity (Feldmann et al. 2012; Narayanan et al. 2012). Adopting this α_{CO} yields a final 3σ upper limit on the molecular gas mass of C21434 of $M_{\text{H}_2} < 1.1 \times 10^{10} M_{\odot}$; however, this assumption will always introduce a source of systematic uncertainty.

We note that the nondetection of C21434 from the 3 mm continuum is far less constraining. Based on dust emissivity from Dunne et al. (2011) with a dust temperature of 25K and $M_{\text{H}_2}/M_{\text{dust}} = 100$, the 3σ gas mass limit from the 3 mm continuum is $M_{\text{H}_2} < 4.5 \times 10^{10} M_{\odot}$, which is significantly less constraining than the CO(2–1) line flux. This places a very weak limit on $\alpha_{\text{CO}} < 19 M_{\odot} / (\text{K km s}^{-1} \text{ pc}^2)$.

2.4. Ancillary Data Sets

In addition to the new ALMA observations presented herein, we compile a literature sample of high-redshift galaxies with measured molecular gas reservoirs, star formation rates, and stellar sizes measured from rest-frame optical *HST* imaging. We include 38 star-forming galaxies from PHIBSS (CO(3–2),

Tacconi et al. 2010, 2013), 97 (38 at $z > 1$) from PHIBSS2 (Tacconi et al. 2018) and five from Daddi et al. (2010) with PdBI observations of CO(2–1). We note that for PHIBSS2 we have no additional information about the uncertainty or methodology(ies) used to measure effective radii (Tacconi et al. 2018). In the absence of information about, e.g., how well-fit these galaxies are by Sérsic profiles, whether they are fit in the rest-frame optical from high-resolution *HST* imaging, and whether sizes are circularized, we exclude the sample from structural comparisons. We include seven galaxies from Decarli et al. (2016) with robust (flag = 0) rest-frame optical *HST* size measurements (van der Wel et al. 2012) and reliable stellar masses ($\log M_{\star, 3D-HST}/M_{\odot} > 9$), correcting to a Milky Way α_{CO} . Three compact star-forming galaxies with CO(1–0) VLA observations and *HST*/WFC3 imaging, are included from Spilker et al. (2016). Sargent et al. (2015), Hayashi et al. (2018), and Rudnick et al. (2017) represent the only other three CO-based constraints on molecular gas in high-redshift quiescent galaxies at $z = 1.4277$, $z = 1.451$, and $z = 1.62$. We include the Sargent et al. (2015) galaxy adopting the Onodera et al. (2012) stellar mass, which also uses a Chabrier (2003) IMF. For maximum consistency, we include 3D-*HST* stellar masses (Skelton et al. 2014) for the Rudnick et al. (2017) galaxies. Gobat et al. (2018) reported a detection of $f_{\text{H}_2} \sim 16\%$ in quiescent galaxies at $z \sim 1.8$, correcting to a Chabrier (2003) IMF using a factor of 0.55 based on Longhetti & Saracco (2009), obtained via a median stacking analysis of mid-IR, far-IR, submillimeter, and radio observations to study the average dust-continuum emission from 977 quiescent galaxies. We adopt the effective radius from the van der Wel et al. (2014) quiescent size–mass relation at the average mass ($\langle M_{\star} \rangle = 6 \times 10^{10} M_{\odot}$) for this sample. At intermediate redshift ($z \sim 0.6$), Suess et al. (2017) present two massive post-starburst galaxies with significant molecular gas reservoirs; however, the stellar sizes are unconstrained for these galaxies. We also include eight quiescent galaxies at $z \sim 0.8$ from Spilker et al. (2018).

Finally we include CO(1–0)-based data from the COLD-GASS survey (Saintonge et al. 2011a, 2011b, 2012) as a low-redshift benchmark. For comparison to $1 < z < 2$ galaxies, we include galaxies from the 3D-*HST* photometric catalogs (Brammer et al. 2011; Skelton et al. 2014), with UV+IR star formation rates (Whitaker et al. 2014) and structural parameters derived from *HST*/WFC3 imaging (van der Wel et al. 2012). We separate the 3D-*HST* galaxies into star-forming and quiescent populations using Whitaker et al. (2012) rest-frame $U - V$ and $V - J$ color cuts.

3. H₂ Reservoirs across the Galaxy Population

Figure 3 shows C21434 (red star) and the literature sample in effective radius and star formation rate as a function of stellar mass, colored by redshift. The vast majority of galaxies with measured gas reservoirs are extended (left) and star-forming (right), consistent with the overall distribution of star-forming galaxies at this redshift. We note that Spilker et al. (2016) targets are selected to be structurally similar to compact quiescent galaxies like C21434. The Sargent et al. (2015) target is sufficiently massive that it overlaps with the star-forming population due to the steep quiescent size–mass relation (e.g., Mowla et al. 2018). The Gobat et al. (2018) sample is clearly the least star-forming $z > 1$ sample, but due to its stacked nature its size is assumed on average. The full sample is heterogeneous in redshift; however, we see the expected trend

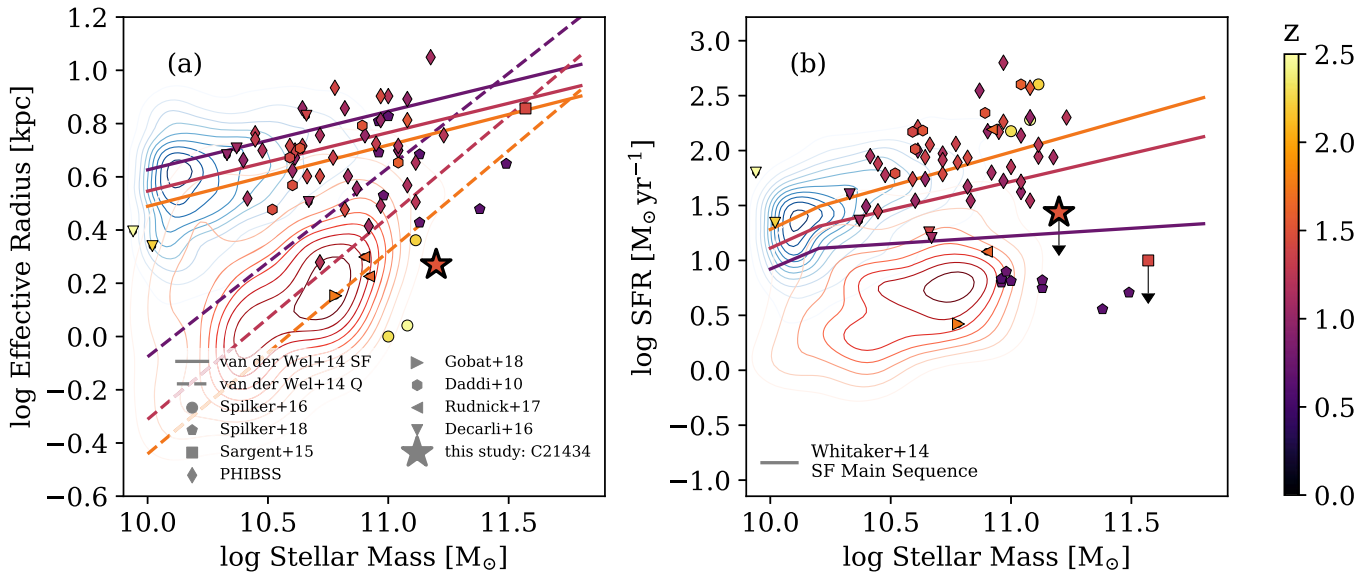


Figure 3. Galaxy size (left panel) and star formation rates (right panel) as a function of stellar mass. Contours indicate the location of star-forming (blue contours) and quiescent (red contours) galaxies at $1 < z < 2$ in the 3D-*HST* survey. van der Wel et al. (2014) size–mass relations at several epochs (left) and Whitaker et al. (2014) star formation “main sequence” (right) are included as solid and dashed lines for star-forming or quiescent relations respectively. Galaxies from a variety of samples with molecular gas and rest-frame optical size measurements are indicated by colored symbols (Daddi et al. 2010; Tacconi et al. 2010; Sargent et al. 2015; Decarli et al. 2016; Spilker et al. 2016, 2018; Rudnick et al. 2017; Gobat et al. 2018), with C21434 indicated by a red star. Symbol and line colors indicate the redshift of each object or relation. The majority of targeted galaxies have been limited to the extended (left) and highly star-forming population (right); however, a few studies have begun to probe the reservoirs of galaxies like C21434 with compact structures (Spilker et al. 2016) and/or minimal ongoing star formation (Sargent et al. 2015; Rudnick et al. 2017; Gobat et al. 2018).

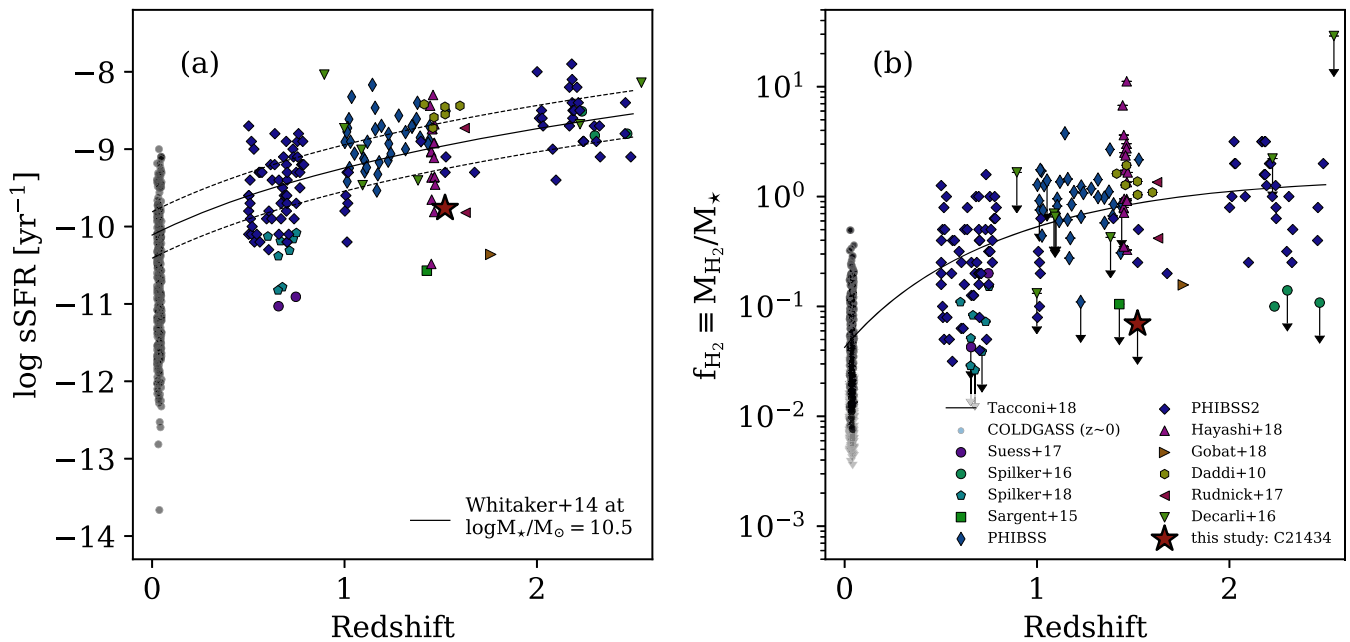


Figure 4. Specific star formation rate (sSFR, left) and molecular gas fraction (f_{H_2} , right) vs. redshift for the high-redshift sample as well as targets from the nearby COLDGASS survey (Saintonge et al. 2011a), which span the full range of galaxy demographics at $z \sim 0$. Expected redshift evolution in sSFR (left, Whitaker et al. 2014) and in f_{H_2} (right, Tacconi et al. 2018) at $\log M_*/M_\odot \sim 10.5$ are indicated by black lines. In the left panel, quiescent targets dramatically stand out from star-forming counterparts in sSFR; however, in gas fractions the Spilker et al. (2016) compact star-forming galaxies have similarly depleted molecular gas reservoirs.

that galaxies at higher- z (lighter symbols) are more compact (Figure 3(a)) and have higher SFRs (Figure 3(b)) at fixed mass.

Locally, there is a strong correlation between molecular gas supply and the efficiency of star formation (e.g., Saintonge et al. 2011a). In Figure 4 we show the specific star formation rate (sSFR \equiv SFR/ M_*) and molecular gas fractions ($f_{\text{H}_2} \equiv M_{\text{H}_2}/M_*$) as a function of redshift. For reference we include the

evolution of sSFR at $M_* \sim 10^{10.5}$ (Whitaker et al. 2014) as a black line in Figure 4(a) and the Tacconi et al. (2018) redshift evolution (with $\log(1+z)^2$ scaling) for a star-forming $M_* \sim 10^{10.5}$ galaxy in Figure 4(b). In this projection, the quiescent galaxies stand out (galaxies with low sSFR in Figure 4(a)), along with the compact star-forming galaxies (yellow circles, Spilker et al. 2016) in Figure 4(b), as deficient

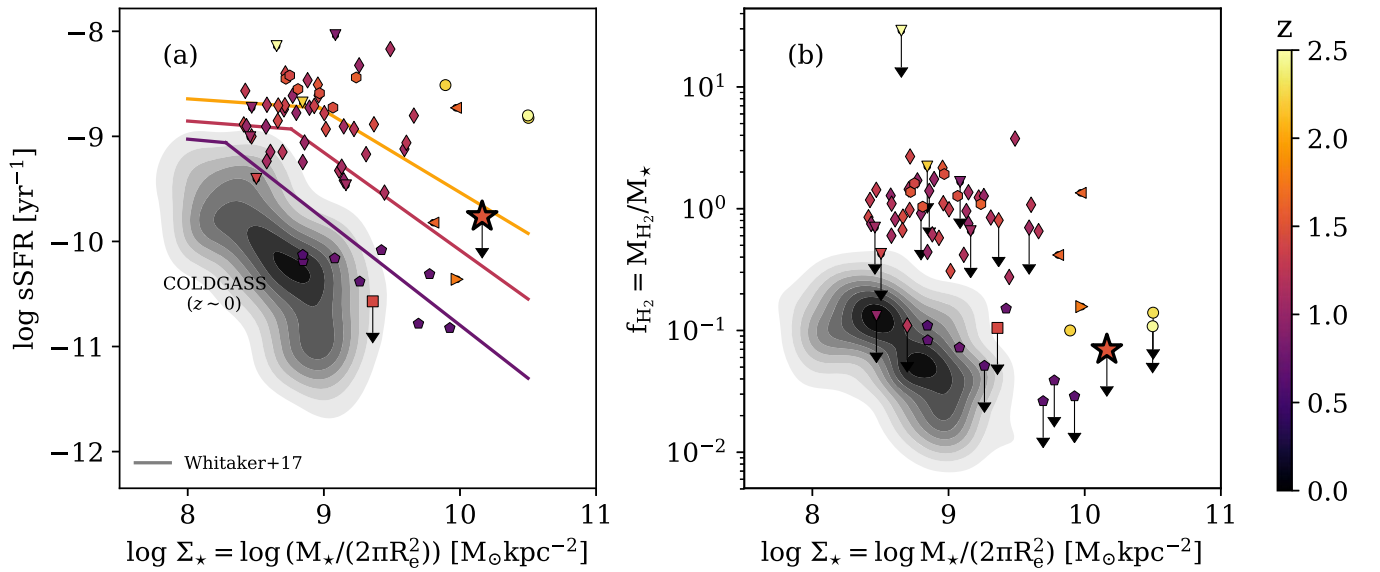


Figure 5. Specific star formation rate and H_2 gas fraction vs. stellar mass surface density for galaxies at $z \sim 0$ (gray contours, COLDGASS survey) and at higher redshift (symbols, colored by redshift). Local galaxies lie offset in sSFR (left) and f_{H_2} (right) at fixed stellar density, as expected. Broken power law relations between sSFR and Σ_* for all galaxies at $\langle z \rangle = 0.75, 1.25,$ and 1.75 from Whitaker et al. (2017) are indicated by solid lines in panel (a). Although some galaxies at fixed density and redshift overlap in sSFR (a), gas fractions at all redshifts drop dramatically for the most dense galaxies ($\Sigma_* \gtrsim 5 \times 10^9 M_\odot \text{ kpc}^{-2}$, b).

in molecular gas. We note that this comparison is relative to coeval galaxies; the molecular gas fractions of all “depleted” galaxies at high-redshift are still consistent with many local gas-rich star-forming galaxies.

The similarity between the molecular gas reservoirs in compact star-forming and quiescent galaxies suggests a connection between stellar structures and gas fractions. This may be expected from the observed correlations between star formation and galaxy structure observed at all redshifts (e.g., Franx et al. 2008; Whitaker et al. 2017, and references therein). Figure 5(a) shows sSFR and f_{H_2} versus stellar mass surface density ($\Sigma_* \equiv M_*/(2\pi R_e^2)$) where galaxies with lower sSFR tend to have higher densities at any epoch. The average relations for all galaxies in the 3D-*HST* survey at several epochs are included from Whitaker et al. (2017). Remarkably, although the high Σ_* populations are still mildly overlapping in sSFR with extended galaxies, they separate relatively cleanly in gas fraction (Figure 5(b)). This trend exists at each epoch such that less dense galaxies also have higher f_{H_2} . We note that although this appears to be at tension with recent results from Freundlich et al. (2019), which found no trends in gas mass with stellar surface density, that study is primarily based on extended star-forming galaxies with much lower stellar densities; above $\Sigma_* \gtrsim 9.5$ those data also show hints of depleted M_{H_2} .

4. Discussion

Measurements of the molecular gas contents of quenching galaxies provide critical constraints on the physics driving future star formation. Our deep ALMA observation provides the most stringent constraint on the molecular gas reservoir of a quiescent galaxy beyond $z > 1$ and one of the deepest outside the local universe to date. This limit of $f_{\text{H}_2} \lesssim 7\%$, indicates that the gas depletion in this galaxy was effective and nearly complete. As discussed in Section 2.4, the few constraints that exist among quenched galaxies beyond $z > 1$ collectively indicate a surprising diversity of molecular gas contents,

ranging from f_{H_2} between $\lesssim 10\%$ (Sargent et al. 2015) to $\sim 16\%$ (Gobat et al. 2018) to as high as $\sim 40\%$ in two cluster galaxies (Rudnick et al. 2017; Hayashi et al. 2018). Similar diversity has been observed among quiescent and post-starburst galaxies at low and intermediate redshifts (4%–30%; French et al. 2015; Suess et al. 2017; Spilker et al. 2018). This object, combined with the two gas-rich post-starburst galaxies (Suess et al. 2017), provides evidence for a similar diversity of gas contents in galaxies at $z > 0$ immediately following quenching. The scatter or distribution of molecular gas content with respect to stellar age may encode critical hints toward the physics driving the quenching process in massive galaxies.

We also find that the molecular gas fraction is strongly correlated with the stellar mass surface density at all epochs, both in the local universe (see also, e.g., Saintonge et al. 2011a) and since $z \sim 2$. α_{CO} remains a source of systematic uncertainty, but we note that, for example, adopting a lower value of α_{CO} for the galaxies with the highest stellar densities (e.g., Bolatto et al. 2013) would only strengthen the break between populations in Figure 5(b). Although idealized merger simulations predict a broad range of α_{CO} (e.g., Narayanan et al. 2011) as could be relevant for compact merger remnants, it would be very difficult to produce higher α_{CO} values unless the galaxies have significantly subsolar metallicity. We expect the dense galaxies in these samples to be metal-rich given their high stellar masses and therefore posit that adopting a Milky Way α_{CO} is a conservative assumption. This result may be another manifestation of the reasonable correlation between stellar and gas structures (e.g., Tacconi et al. 2013) and the Kennicutt–Schmidt relation (Schmidt 1959; Kennicutt 1998). Perhaps this implies that high stellar densities facilitate efficient star formation and gas consumption. However, we note that such correlations may also result from the effects of progenitor biases (Lilly & Carollo 2016). Alternatively, more compact galaxies have higher stellar velocity dispersions and therefore likely host larger supermassive black holes, which may in turn drive stronger feedback (e.g., Magorrian et al. 1998). There is a well-established correlation between galaxy density—either on

average or within a central region—and stellar populations and ongoing star formation (e.g., Kauffmann et al. 2003; Franx et al. 2008; Cheung et al. 2012; Fang et al. 2013; Whitaker et al. 2017, and references therein). The correlations in Figure 5(b) suggest that this may be related to the gas reservoirs fueling that star formation, but given the sparse sampling of very dense galaxies at high redshift we are limited to speculation. It is also possible that causality points in the other direction and that the sharp transition in gas fraction is driven by differences in specific star formation rates, which in turn correlate with structures. We stress that obtaining larger samples will be crucial in quantifying that diversity and exploring correlations with stellar populations and structures, particularly pushing to high redshift where observations probe closer to the quenching epoch for massive galaxies.

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References

Bezanson, R., van Dokkum, P., van de Sande, J., Franx, M., & Kriek, M. 2013, *ApJL*, 764, L8
 Bolatto, A. D., Wolfire, M., & Leroy, A. K. 2013, *ARA&A*, 51, 207
 Brammer, G. B., Whitaker, K. E., van Dokkum, P. G., et al. 2011, *ApJ*, 739, 24
 Bruzual, G., & Charlot, S. 2003, *MNRAS*, 344, 1000
 Cappellari, M., & Emsellem, E. 2004, *PASP*, 116, 138
 Chabrier, G. 2003, *PASP*, 115, 763
 Cheung, E., Faber, S. M., Koo, D. C., et al. 2012, *ApJ*, 760, 131

Daddi, E., Bournaud, F., Walter, F., et al. 2010, *ApJ*, 713, 686
 Davidzon, I., Ilbert, O., Laigle, C., et al. 2017, *A&A*, 605, A70
 Davis, T. A., Alatalo, K., Bureau, M., et al. 2013, *MNRAS*, 429, 534
 Davis, T. A., Alatalo, K., Sarzi, M., et al. 2011, *MNRAS*, 417, 882
 Davis, T. A., Greene, J., Ma, C.-P., et al. 2016, *MNRAS*, 455, 214
 Decarli, R., Walter, F., Aravena, M., et al. 2016, *ApJ*, 833, 70
 Dunne, L., Gomez, H. L., da Cunha, E., et al. 2011, *MNRAS*, 417, 1510
 Fang, J. J., Faber, S. M., Koo, D. C., & Dekel, A. 2013, *ApJ*, 776, 63
 Feldmann, R., Gnedin, N. Y., & Kravtsov, A. V. 2012, *ApJ*, 758, 127
 Franx, M., van Dokkum, P. G., Schreiber, N. M. F., et al. 2008, *ApJ*, 688, 770
 French, K. D., Yang, Y., Zabludoff, A., et al. 2015, *ApJ*, 801, 1
 French, K. D., Zabludoff, A. I., Yoon, I., et al. 2018, *ApJ*, 861, 123
 Freundlich, J., Combes, F., Tacconi, L. J., et al. 2019, *A&A*, 622, A105
 Glazebrook, K., Schreiber, C., Labbé, I., et al. 2017, *Natur*, 544, 71
 Gobat, R., Daddi, E., Magdis, G., et al. 2018, *NatAs*, 2, 239
 Hayashi, M., Tadaki, K.-i., Kodama, T., et al. 2018, *ApJ*, 856, 118
 Hunt, Q., Bezanson, R., Greene, J. E., et al. 2018, *ApJL*, 860, L18
 Kauffmann, G., Heckman, T. M., White, S. D. M., et al. 2003, *MNRAS*, 341, 54
 Kennicutt, R. C. 1998, *ARA&A*, 36, 189
 Kriek, M., van Dokkum, P. G., Labbé, I., et al. 2009, *ApJ*, 700, 221
 Lilly, S. J., & Carollo, C. M. 2016, *ApJ*, 833, 1
 Longhetti, M., & Saracco, P. 2009, *MNRAS*, 394, 774
 Magorrian, J., Tremaine, S., Richstone, D., et al. 1998, *AJ*, 115, 2285
 Man, A., & Belli, S. 2018, *NatAs*, 2, 695
 Mowla, L., van Dokkum, P., Brammer, G., et al. 2018, arXiv:1808.04379
 Muzzin, A., Marchesini, D., Stefanon, M., et al. 2013, *ApJ*, 777, 18
 Muzzin, A., Marchesini, D., van Dokkum, P. G., et al. 2009, *ApJ*, 701, 1839
 Narayanan, D., Krumholz, M., Ostriker, E. C., & Hernquist, L. 2011, *MNRAS*, 418, 664
 Narayanan, D., Krumholz, M. R., Ostriker, E. C., & Hernquist, L. 2012, *MNRAS*, 421, 3127
 Onodera, M., Renzini, A., Carollo, M., et al. 2012, *ApJ*, 755, 26
 Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2002, *AJ*, 124, 266
 Rudnick, G., Hodge, J., Walter, F., et al. 2017, *ApJ*, 849, 27
 Saintonge, A., Kauffmann, G., Kramer, C., et al. 2011a, *MNRAS*, 415, 32
 Saintonge, A., Kauffmann, G., Wang, J., et al. 2011b, *MNRAS*, 415, 61
 Saintonge, A., Tacconi, L. J., Fabello, S., et al. 2012, *ApJ*, 758, 73
 Sargent, M. T., Daddi, E., Bournaud, F., et al. 2015, *ApJL*, 806, L20
 Schmidt, M. 1959, *ApJ*, 129, 243
 Sérsic, J. L. 1968, Atlas de Galaxias Australes (Cordoba, Argentina: Observatorio Astronomico)
 Skelton, R. E., Whitaker, K. E., Momcheva, I. G., et al. 2014, *ApJS*, 214, 24
 Spilker, J., Bezanson, R., Barišić, I., et al. 2018, *ApJ*, 860, 103
 Spilker, J. S., Bezanson, R., Marrone, D. P., et al. 2016, *ApJ*, 832, 19
 Straatman, C. M. S., Labbé, I., Spitler, L. R., et al. 2014, *ApJL*, 783, L14
 Suess, K. A., Bezanson, R., Spilker, J. S., et al. 2017, *ApJL*, 846, L14
 Tacconi, L. J., Genzel, R., Neri, R., et al. 2010, *Natur*, 463, 781
 Tacconi, L. J., Genzel, R., Saintonge, A., et al. 2018, *ApJ*, 853, 179
 Tacconi, L. J., Neri, R., Genzel, R., et al. 2013, *ApJ*, 768, 74
 Tomczak, A. R., Quadri, R. F., Tran, K.-V. H., et al. 2014, *ApJ*, 783, 85
 van der Wel, A., Bell, E. F., Häussler, B., et al. 2012, *ApJS*, 203, 24
 van der Wel, A., Franx, M., van Dokkum, P. G., et al. 2014, *ApJ*, 788, 28
 Vazdekis, A. 1999, *ApJ*, 513, 224
 Whitaker, K. E., Bezanson, R., van Dokkum, P. G., et al. 2017, *ApJ*, 838, 19
 Whitaker, K. E., Franx, M., Leja, J., et al. 2014, *ApJ*, 795, 104
 Whitaker, K. E., Kriek, M., van Dokkum, P. G., et al. 2012, *ApJ*, 745, 179
 Whitaker, K. E., Labbe, I., van Dokkum, P. G., et al. 2011, *ApJ*, 735, 86
 Young, L. M., Bureau, M., Davis, T. A., et al. 2011, *MNRAS*, 414, 940
 Young, L. M., Scott, N., Serra, P., et al. 2014, *MNRAS*, 444, 3408